

**ALMAHATA SITTA NEWS: WELL-KNOWN VARIETIES AND NEW SPECIES IN THE ZOO.**A. Bischoff<sup>1</sup>, S. Ebert<sup>1</sup>, M. Patzek<sup>1</sup>, M. Horstmann<sup>1</sup>, A. Pack<sup>2</sup>, and S. Decker<sup>3</sup>.<sup>1</sup>Institut für Planetologie, Wilhelm-Klemm-Str. 10, D-48149 Münster, Germany. E-mail: bischoa@uni-muenster.de.<sup>2</sup>Uni Göttingen, Geowissenschaftliches Zentrum, Goldschmidtstr. 1, D-37077 Göttingen, Germany.<sup>3</sup>Meteorite-Museum, Oberstr. 10a, D-55430 Oberwesel, Germany.

**Introduction:** After asteroid 2008 TC<sub>3</sub> impacted Earth in 2008, a highly diverse range of meteorite types was identified among the numerous meteorite fragments collected in the Almahata Sitta strewn field, including various types of ureilitic rocks and chondrites (e.g. [1-10]). Here, mineralogical characteristics of 18 new samples (MS-MU-021 – MS-MU-038) are presented, complemented with the oxygen isotope signature of MS-MU-021..

**Results:** The samples represent (like earlier studied fragments) one single lithology weighing between 7.4 (MS-MU-028) and 221.5 g (MS-MU-034). Together with MS-MU-034 some other large samples of the strewn field are characterized: MS-MU-036 (177.1 g), -037 (149.2 g), and -038 (148.2 g). Among the samples are 5 E chondrites, 12 samples of ureilitic origin (including a trachyandesite (MS-MU-035)), and an enstatite- and metal-rich achondrite.

*E-chondrites:* Based on the low Si-concentration in metal (0.8-1.3 wt%) all five enstatite chondrites are from the EL-subgroup (2 EL3, 1 EL4, 2 EL6 (MS-MU-024, -026)). The EL3 chondrites (MS-MU-023, -031) can be characterized by their well-preserved chondritic texture and the occurrence of olivine. Especially in MS-MU-031 abundant Fa-poor olivine was found within the porphyritic chondrules. Olivine is lacking in MS-MU-029 (EL4).

*Ureilites:* MS-MU-022, -034, -037, and -038 are coarse-grained ureilites. Within MS-MU-022 olivine has cores of Fa<sub>1-13</sub>, whereas the Fa-contents of olivine cores from MS-MU-034 and -038 are much higher and very similar (23-24 mol%). Olivine within MS-MU-034 contains abundant tiny Cr-rich inclusions (exsolved chromites; <<1 μm). MS-MU-037 has abundant pyroxene (~Fs<sub>18</sub>Wo<sub>5</sub>) and broad reduction rims around the olivine (Fa<sub>19-21</sub>). Two ureilitic fragments have complex textures with highly variable grain sizes (MS-MU-021 (Fig. 1), -028). The composition of the olivine cores in MS-MU-028 is Fa<sub>15-17</sub>, whereas Fa-contents of up to 26 mol% were analyzed in olivines from MS-MU-021. In this sample a small number of large olivine grains are embedded in a fine-grained groundmass (Fig. 1). The bulk sample has δ<sup>17</sup>O = 3.98‰ and δ<sup>18</sup>O = 8.29‰, which is consistent with an ureilitic origin.

MS-MU-025, -027, -030, -032, and -033 are heavily-shocked, fine-grained ureilites. MS-MU-027 shows a high abundance of C-phase(s) and typical olivine has very low Fa contents (<3 mol%). The samples MS-MU-030 and MS-MU-032 contain abundant metal and sulfide and their core olivines have about 10 and 9-12 mol% Fa, respectively. MS-MU-033 is a typical fine-grained ureilite with olivine cores of Fa<sub>19-21</sub>.

*Ureilitic trachyandesite:* The thin section inspection revealed that this sample (MS-MU-035) is similar to MS-MU-011 (ALM-A [9]) having abundant (~65 vol%) subhedral, zoned anorthoclase and/or plagioclase (typically ~An<sub>9-25</sub>) embedding Cr-bearing Ca-pyroxene (~Fs<sub>20</sub>Wo<sub>38</sub>) and Ca-poor pyroxene (~Fs<sub>35.5</sub>Wo<sub>8</sub>) (Fig. 2).

*Enstatite- and metal-rich achondrite:* MS-MU-036 is an unusual enstatite- and metal-rich achondrite with three coexisting pyroxenes (~En<sub>98.5</sub>Wo<sub>1.3</sub>, ~En<sub>96.5</sub>Wo<sub>3.2</sub>, and ~En<sub>60</sub>Wo<sub>40</sub>). The silicates are embedded within a Si-bearing metal matrix (Fig. 3). Additional phases include oldhamite, alabandite, and daubréelite. Based on texture and mineralogy, it might represent a unique type of meteorite, however, similarities exist to MS-MU-019 and Itqiy [11].

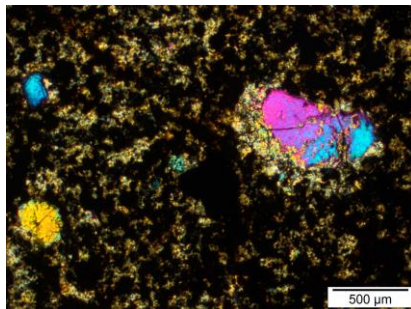


Fig. 1: MS-MU-021

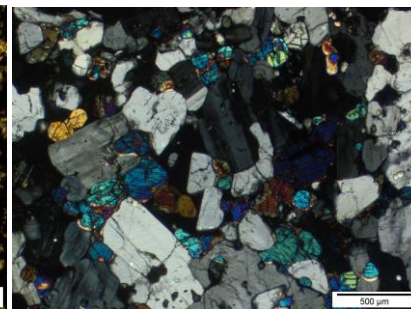


Fig. 2: MS-MU-035

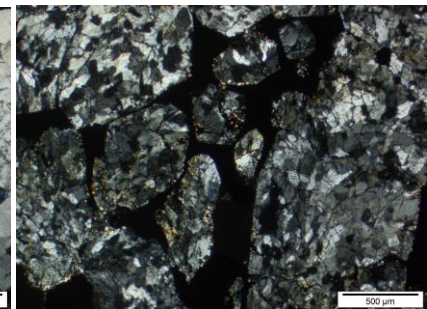


Fig. 3: MS-MU-036 (metal: black)

**References:** [1] Jenniskens P. et al. 2009. *Nature* 458:485-488. [2] Bischoff A. et al. 2010. *Meteoritics & Planet. Sci.* 45:1638-1656. [3] Horstmann M. and Bischoff A. 2014. *Chemie der Erde - Geochemistry* 74:149-184. [4] Goodrich C. A. et al. 2014. *Elements* 10:31-37 [5] Bischoff A. et al. 2012. *Meteoritics & Planet. Sci.* 47:A71. [6] Horstmann M. et al. 2010. *Meteoritics & Planet. Sci.* 45:1657-1667. [7] Zolensky M. E. et al. 2010. *Meteoritics & Planet. Sci.* 45:1618-1637. [8] Horstmann M. et al. 2012. *Meteoritics & Planet. Sci.* 47:A193. [9] Bischoff A. et al. 2014. *Proceedings National Academy Sciences* 111:12689-12692. [10] Bischoff A. et al. 2015. *Meteoritics & Planet. Sci.* 50:#5092. [11] Patzer A. et al. (2001) *Meteoritics & Planet. Sci.* 36:1495-1505.