

ERROR ANALYSIS OF LUNAR BOULDERS SIZES

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Introduction

- Boulder size-frequency distributions (SFD) play a key role in understanding small-scale erosion processes and the rate at which rock become regolith [1,2].
- Lunar Reconnaissance Orbiter (LRO) Narrow Angle Camera (NAC) images (resolution: 0.5-1 m/pixel, ~60°-80° incidence angle) are used to map boulders.

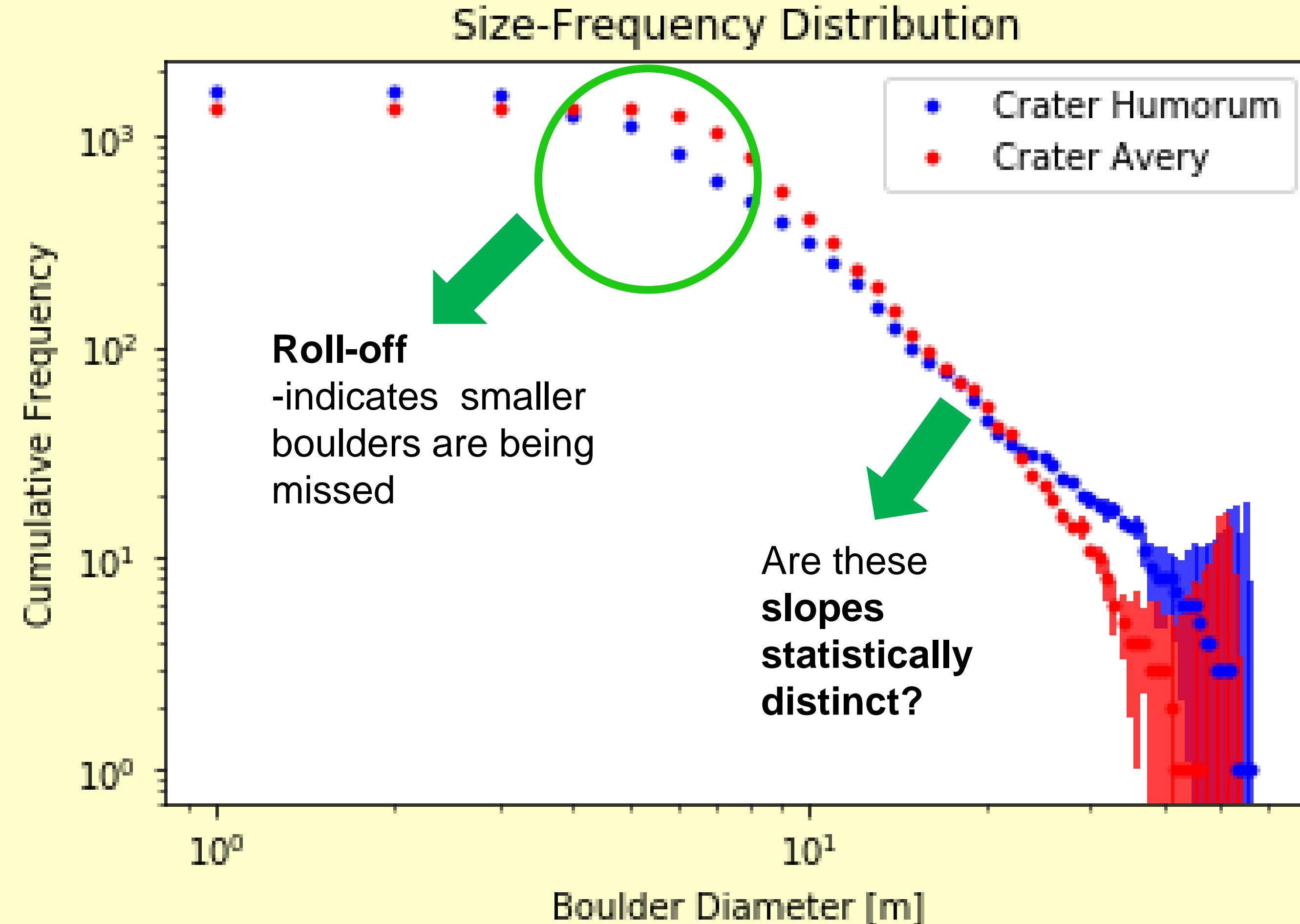


Fig.2: Cumulative size-frequency distributions of boulders around Avery and Humorum craters.

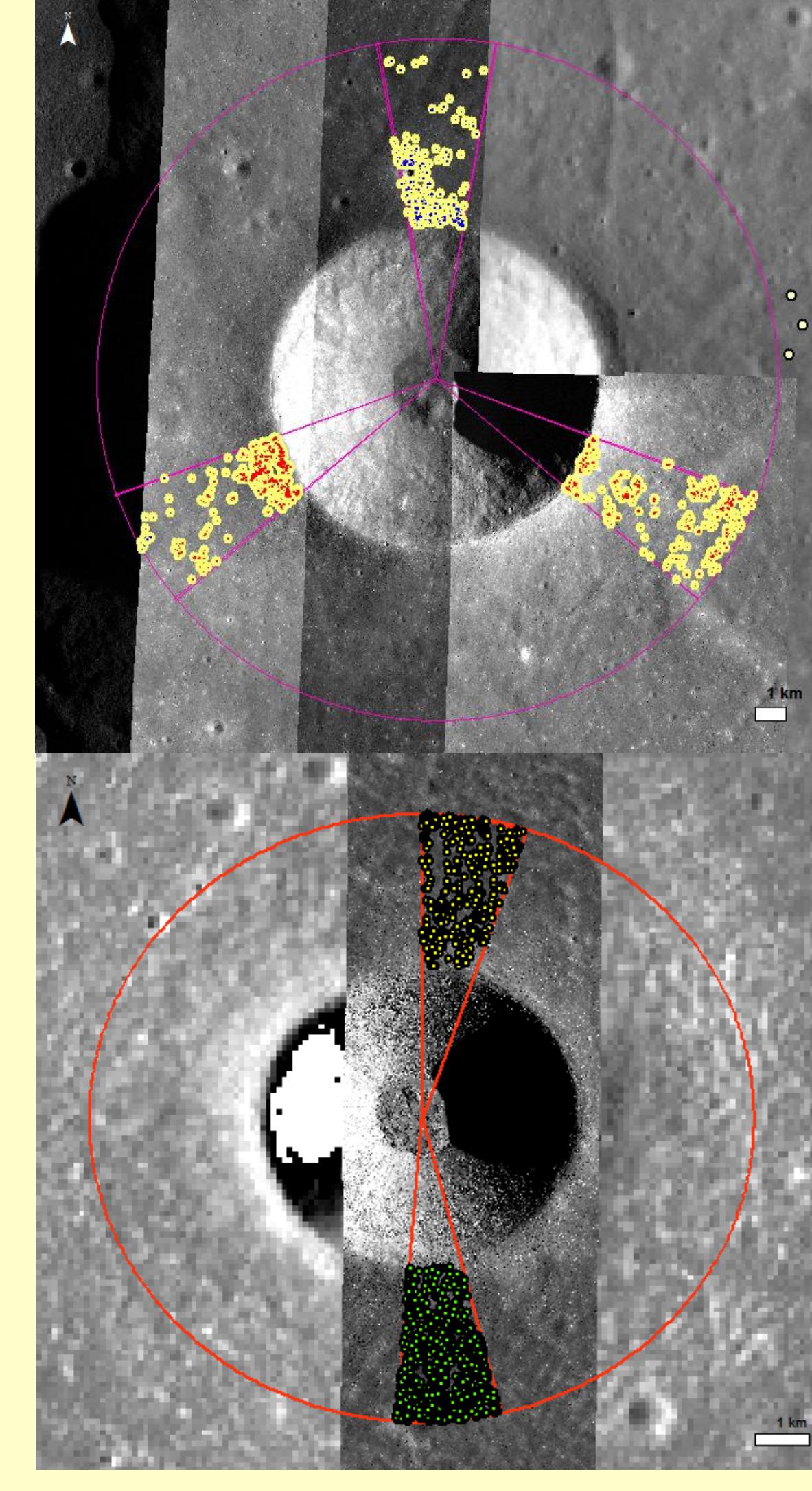


Fig.1: Boulder counts 1 radii out from the crater rim. Top: Avery crater, 1.4°S, 81.4°E, diameter = 9 km, NACs M1123064231, M1156044386, M1924474. Bottom: Crater in Mare Humorum, 319.207°E, -24.051°N, diameter = 5.6 km, NACs M114220623.

- Slope of cumulative SFD is of interest → varies for rock populations that have experienced different levels of comminution [3].

Objectives:

- Determine smallest boulder size that can be mapped with confidence
- Quantify the effect of size uncertainties for all boulder diameters

Methods

- Ellipse fitting tool in ArcMap CraterTools is used to map boulders. Semi-major and minor axes are recorded; the height is assumed to be equivalent to the minor axis.
- Use rectangle fitted inside an ellipse to determine uncertainty in boulder diameters.
- Volume of ellipsoid is larger than volume of rectangular prism which leads to an overestimation of sizes. The uncertainty comes from this volume discrepancy.
- Our results give an upper bound on the boulder size error.

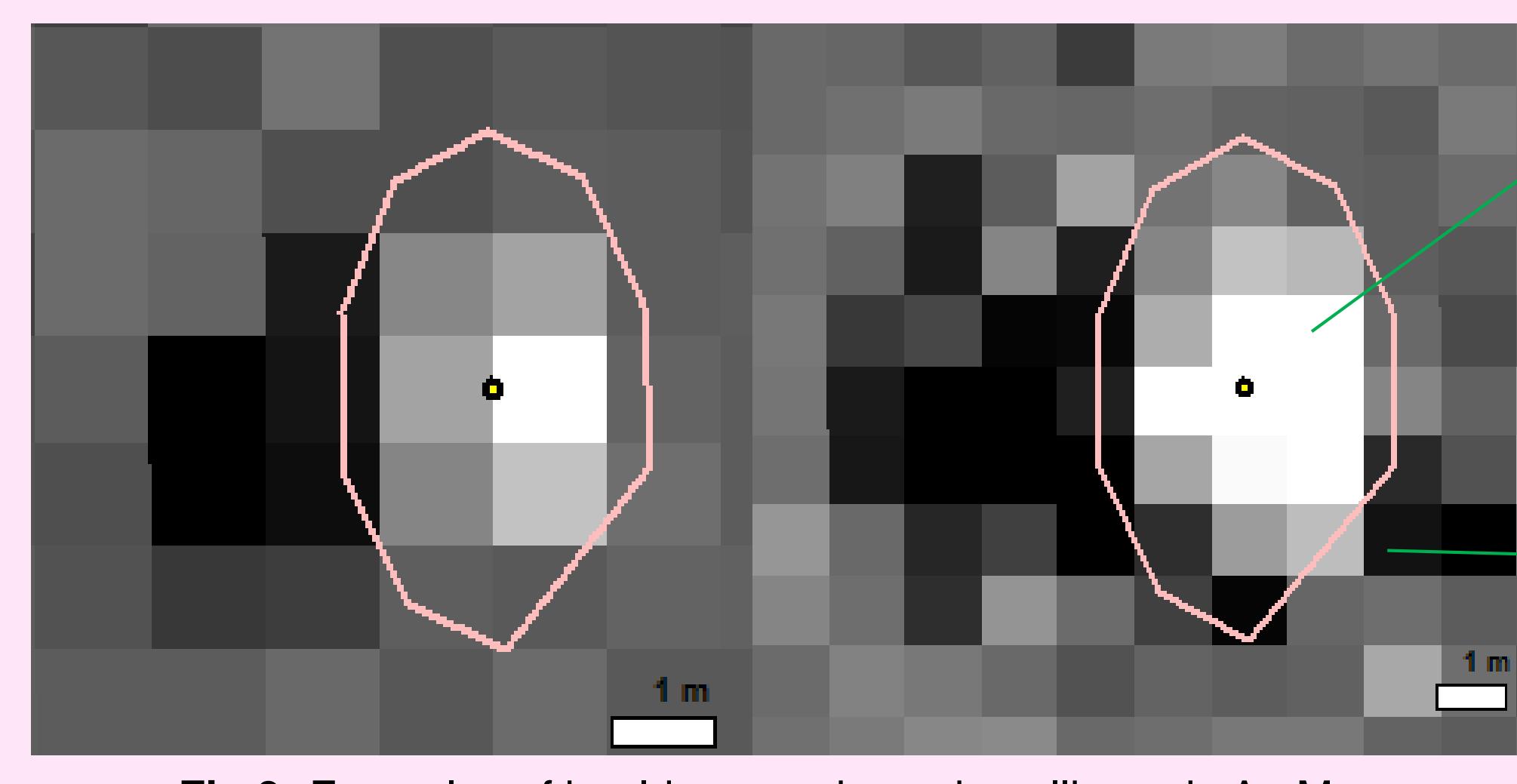


Fig.3: Examples of boulder mapping using ellipses in ArcMap.

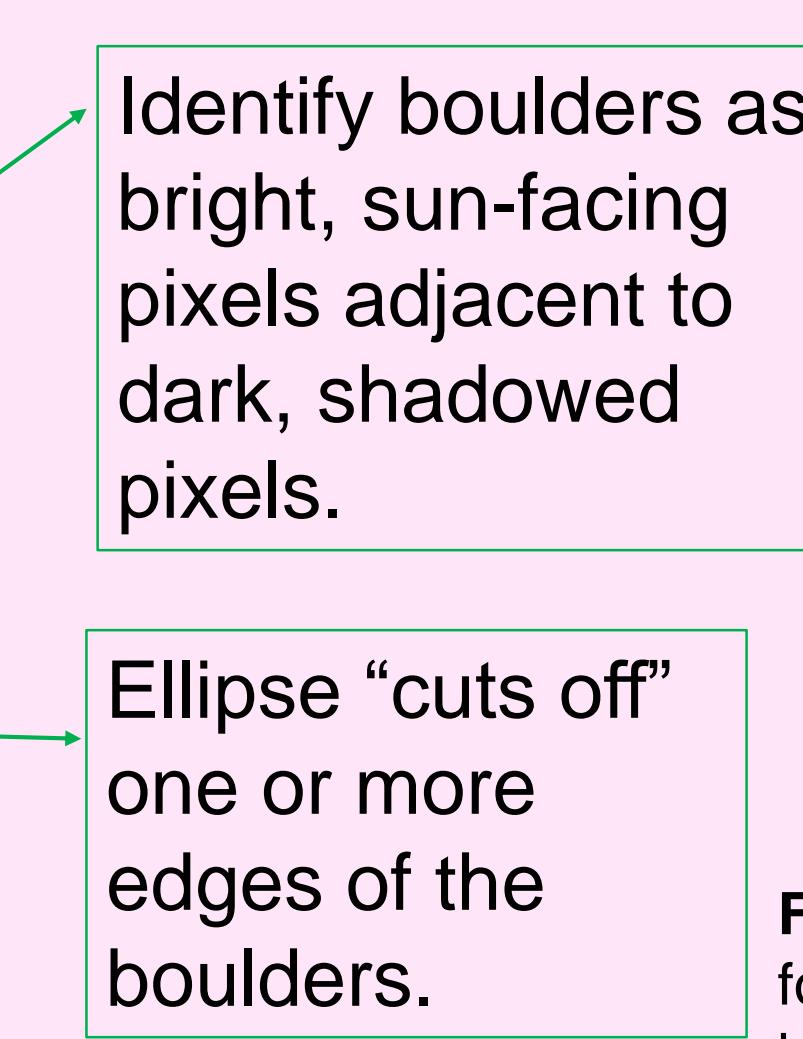


Fig.4: Rectangle fitted inside an ellipse for different pixel arrangements, for a boulder defined by 6 pixels.

Results

- Calculations show an error range of 14-36% on boulder sizes.
- Error calculations performed for boulders with diameter ~2-5 m (2-10 pixels).
- This can be extended to larger boulders because uncertainty is independent of size (depends on rectangle-to-ellipse ratio).

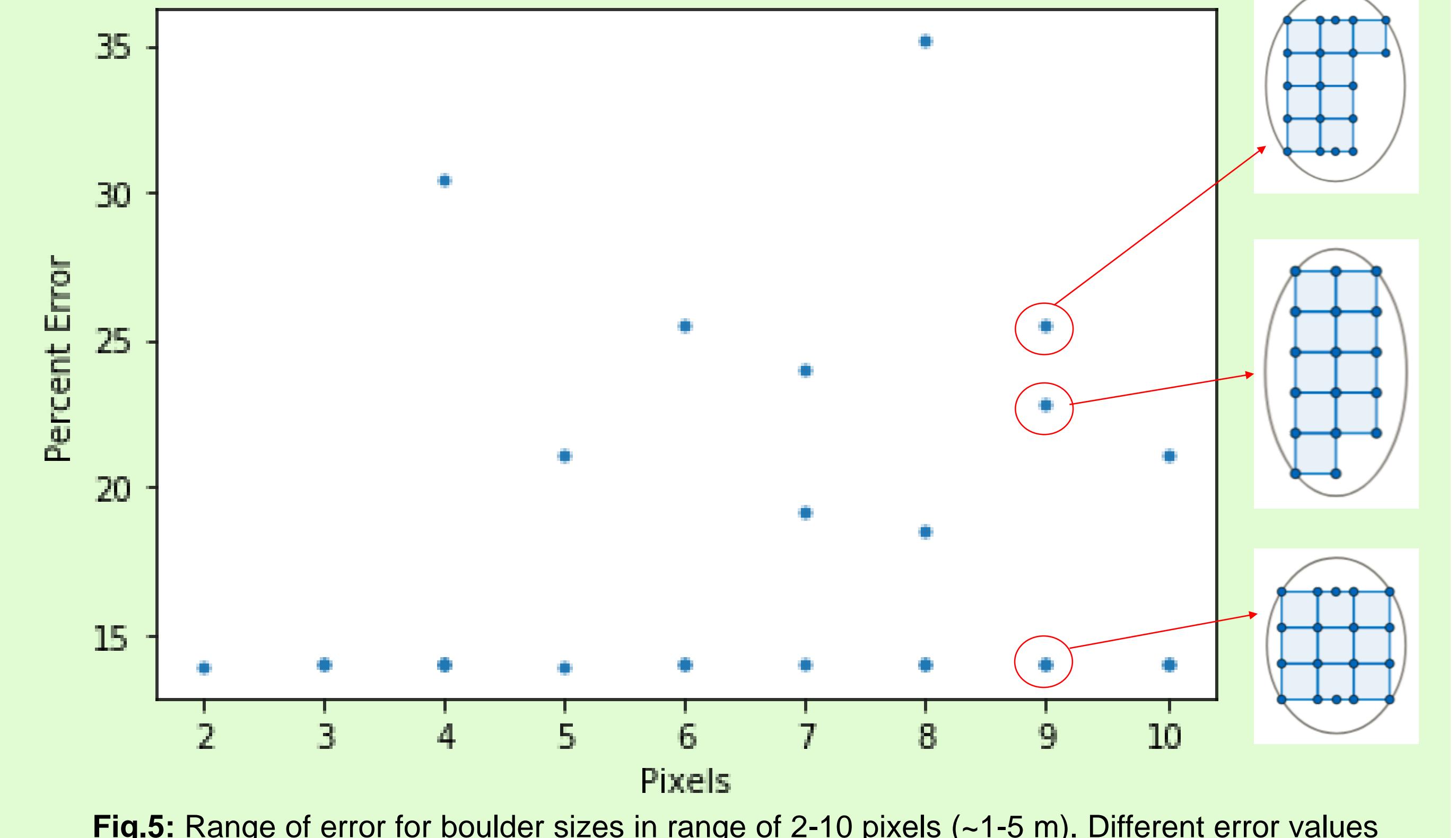


Fig.5: Range of error for boulder sizes in range of 2-10 pixels (~1-5 m). Different error values for a given pixel number is a result of different pixel arrangements as shown to the right.

Testing distinctness of slope

- Assume a normal error distribution and randomly sample from that distribution for each point in the original data, for multiple runs. Take μ = original boulder size and σ = uncertainty on size (10-40%).
- Slope of the linear range in SFD is calculated for each run, in order to compare to original slope.
- For errors in range of 30-40%, the distribution of slope spreads out more compared to the lower uncertainty values.

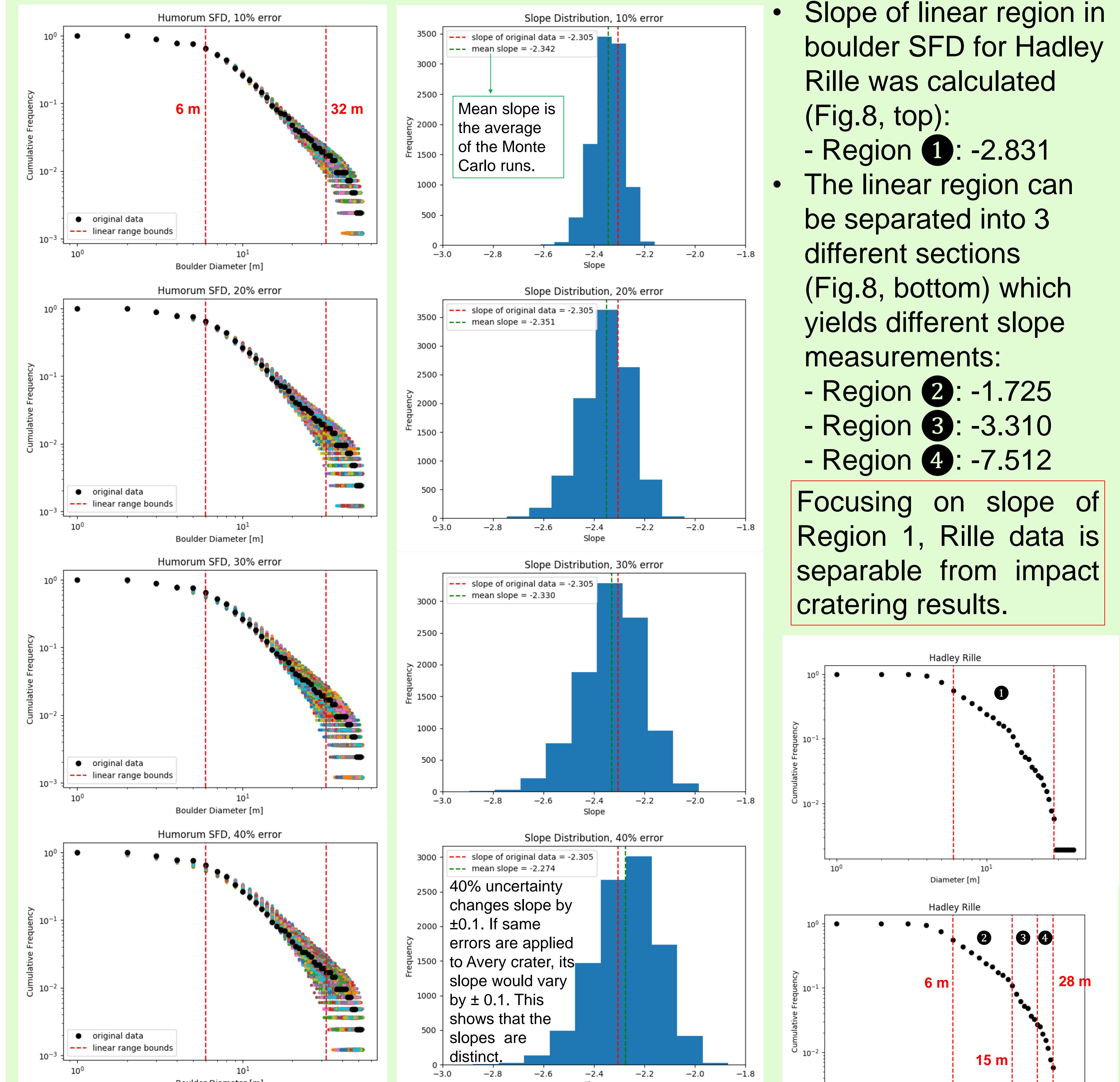


Fig.6: SFDs of Humorum boulders after applying 10-40% error (top to bottom) to each boulder size. The black points represent the original data, and each set of colored points represent a different sampling run (out of 10,000 runs).

Fig.7: Distribution of slope of linear range calculated for each run (total 10,000 runs) for different size uncertainties applied (10-40% top to bottom).

- Slope of linear region in boulder SFD for Hadley Rille was calculated (Fig.8, top):
 - Region 1: -2.831
- The linear region can be separated into 3 different sections (Fig.8, bottom) which yields different slope measurements:
 - Region 2: -1.725
 - Region 3: -3.310
 - Region 4: -7.512

Focusing on slope of Region 1, Rille data is separable from impact cratering results.

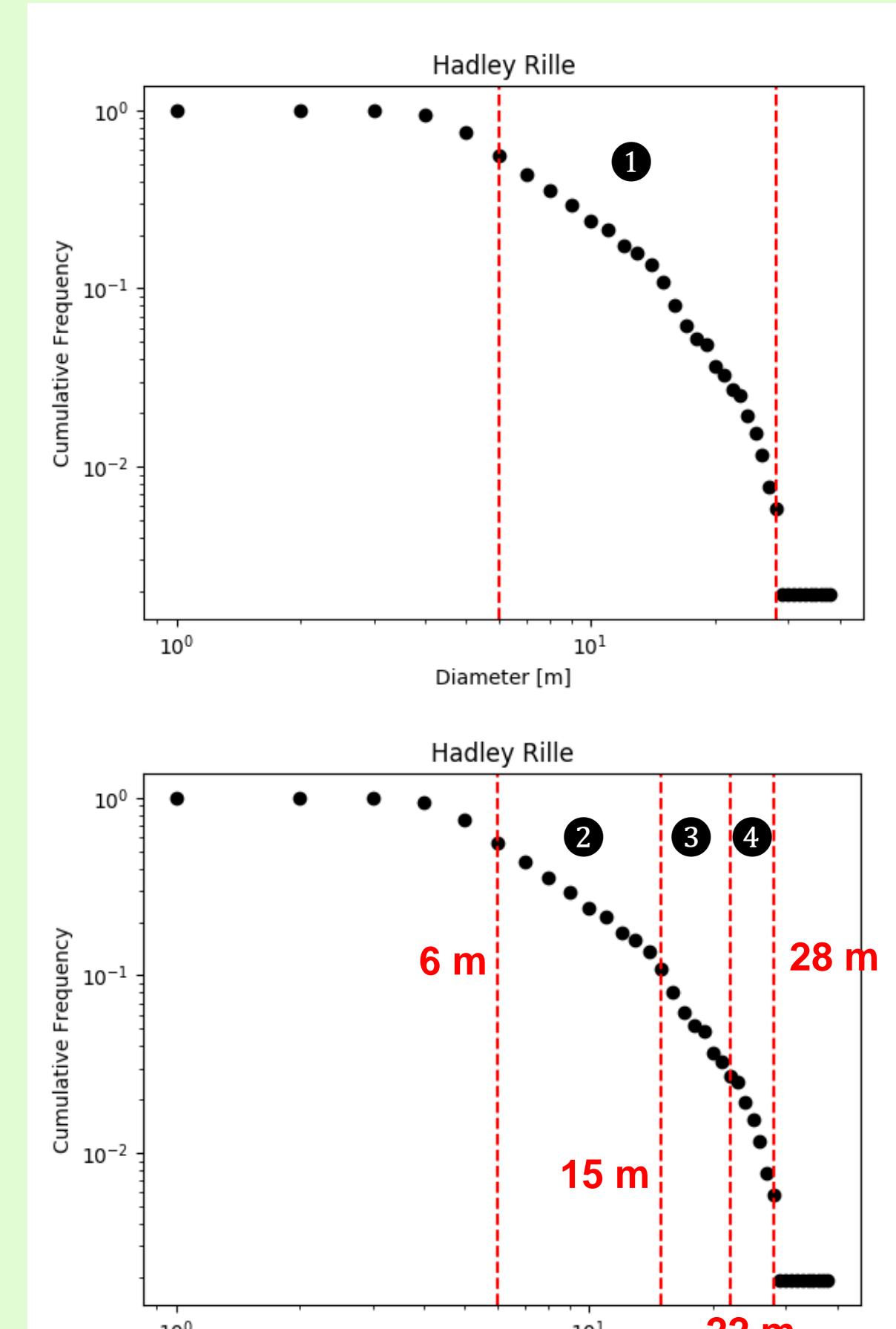


Fig.8: Cumulative SFD of Hadley Rille, 26.69°N, 4.12°E. Data courtesy of Jeff Plescia, JHU/APL.

Conclusions and Future Work

- Monte Carlo runs show a small difference between original and mean slopes.
- Avery crater (slope = -3.120) is different from Humorum crater (slope = -2.305).
- Uncertainties on boulder sizes are one-sided, therefore we need to sample from a distribution other than a normal/Gaussian distribution (e.g. lognormal distribution).

Acknowledgements & References

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References: [1] Gault et al. (1972) *LSL III*, 2713-2744. [2] Hörz et al. (1975) *Moon*, 13, 235-258. [3] Hartmann (1969) *Icarus*, 10, 201-213. [4] Melosh (1984) *Icarus*, 59, 234-260. [5] Ghent et al. (2014) *Geology*, 42, 1059-1062. [6] Li et al. (2017) *PSS*, 146, 30-39. [7] Shoemaker (1965) *JPL Tech. Report #32-700*, 76-134. [8] Arvidson (1979) *Icarus*, 37, 467-474. [9] Watkins et al. (2017) *LPS XLVIII*, Abstract #1245. [10] Bart and Melosh (2010), *Icarus*, 209, 337-357. [11] Demidov and Basilevsky (2014) *Sol. Syst. Res.*, 48, 324-329.