

FEATURES OF THE NORTHERN SMOOTH PLAINS OF MERCURY REVEALED BY DETRENDED MLA TOPOGRAPHY: COMPARISON WITH THE MOON. M. A. Kreslavsky^{1,2}, J. W. Head³, G. A. Neumann⁴, M. T. Zuber⁵, D. E. Smith⁵, ¹Earth and Planetary Sciences, University of California, Santa Cruz, CA, 95064, USA, mkreslav@ucsc.edu; ²MEsLab, Moscow State University of Geodesy and Cartography (MII-GAiK), Moscow, Russia; ³Earth, Environmental and Planetary Sciences, Brown University, Providence, RI, 02912, USA; ⁴Solar System Exploration Division, NASA Goddard Space Flight Center, Greenbelt, MD, USA; ⁵Earth, Atmospheric and Planetary Sciences, MIT, Cambridge, MA, USA.

Introduction: Despite great advances in computer data processing, visual inspection remains a basic analysis tool in geomorphology. Laser altimeter data are, in a sense, too precise for visual perception. For example, the internal ranging precision of the Mercury Laser Altimeter (MLA) [1] onboard the MErcury, Sur-face, Space, ENvironment, GEOchemistry, and Ranging (MESSENGER) mission is about 12 cm, and the topog-raphical amplitudes on Mercury exceed ~5 km, which gives an impressive dynamic range; without special ef-fort this accuracy of the data could remain unused. To fully actualize the great potential of laser altimeter data in sensing low-amplitude gently sloping topographic features we apply a special detrending algorithm, which enables visualization of subtle topographic features in smooth terrains. The same detrending algorithm has been successfully applied to martian [2] and lunar [3] topography.

Detrending procedure: As source data for the de-trended topographic map of Mercury (Fig. 1), we use a preliminary version of the MLA Gridded Data Records from the PDS. Because of MESSENGER orbit con-straints, MLA data coverage is sufficient for detailed topographic mapping only in the northern polar area, therefore we used the interpolated gridded topographic map in polar azimuthal stereographic projection at 500 m per pixel sampling. Northward from ~ 60° latitude the majority of pixels in this source map are obtained from actual MLA ranging measurements rather than by inter-polation.

The detrended topography at each pixel was calcu-lated as a difference between the actual elevation at this pixel and the median elevation of all pixels within a cir-cular window centered at this pixel. The window size defines the spatial scale of the features that are consid-ered as a global trend and thus are subject to removal. We used a window radius of 10 pixels, which is equiva-lent to 10 km in diameter. The reduced resolution ver-sion of the resulting map is shown in Fig. 1.

The filtering algorithm used is similar to high-pass linear filters, however, we use the median in a sliding window to calculate the trend surface, while the linear filters use the weighted arithmetic mean. Any reason-able detrending algorithm, including both our median-based and the traditional linear filters, filters out topog-raphic features much larger than the window, preserves

topographic features much smaller than the window, and distorts features comparable to the window in size. Our algorithm produces less distortion than linear fil-ters, which makes the results of the median-based de-trending much better for visual perception than the re-sults of a linear high-pass filter. This was the principal reason for our choice of the median-based detrending.

The actual accuracy of MLA topographic measure-ments is lower than the internal ranging precision due to inaccuracies in MESSENGER orbit and attitude knowl-edge [1]. Detrending reveals spurious features appear-ing as sub-parallel lines (Fig. 1); they are abundant at lower latitudes, where the data density is lower.

Any detrending algorithm significantly distorts fea-tures comparable to the window size. Due to the non-linear nature of the median-based detrending procedure, distortions of window-scale features are more compli-cated, than for a linear high-pass filter. Typical heavily cratered terrains on Mercury, including vast intercrater plains, have many craters and other topographic fea-tures in the “harmful” 5- 20 km size range. As a result, the detrended topography technique is much more use-ful in the regionally flatter smooth plains [4-6]. Fortu-nately, a significant part of the mapped northern polar area of Mercury (Fig. 1) is occupied by smooth plains of Borealis Planitia.

Topographic texture of smooth plains: Smooth plains in the map (Fig. 1) are covered with a dense quite isotropic network of bright (positive detrended topogra-phy) lineaments that represent wrinkle ridges and smooth convex slope breaks. Almost all these linea-ments are distinguishable in low-sun images taken under favorable illumination geometry. However, the de-trended topography map is better for their detection: it is free of biases related to sun elevation and azimuth.

Many bright lineaments in Borealis Planitia (Fig. 1) are circles; some of them are concentric; many such cir-cles are slope breaks outlining smooth shallow circular depressions, “ghost craters” [4,6,7]. All these circles obviously outline buried craters and double-ring basins that existed before the final emplacement of smooth-plains-forming lavas. On the Moon only a few features of this kind are known [3]. Their abundance on Mercury may be related to the lack of terrain resetting by large basin formation in the northern plains, and by the very

high resurfacing rates of volcanic plains emplacement [4,6].

On the Moon, the detrended topography maps [3] readily reveal smooth low-topography volcanic constructs (mare domes and small shields) amid mare-forming volcanic plains. No similar features are observed in Borealis Planitia [4] on Mercury, which is another aspect of the difference between the volcanic styles of plain-forming volcanism on the Moon [9] and Mercury [4-6].

References: [1] J. Cavanaugh et al. (2007) *Space Sci. Rev.* 131, 451. [2] M. Kreslavsky & J. Head (2001)

JGR 107, 5121, DOI: 10.1029/2001JE001831. [3] M. Kreslavsky et al. (2016) *LPSC* 47, #1331. [4] J. Head et al. (2011) *Science* 333, 1853. [5] B. Denevi et al. (2013) *JGR* 118, 891. [6] L. Ostrach et al. (2015) *Icarus* 250, 602. [7] C. Klimczak et al. (2012) *JGR* 117, E00L03. [8] J. Head & A. Gifford (1980) *Moon* 22, 235. [9] J. Head and L. Wilson (2016) *LPSC* 47, #1189.

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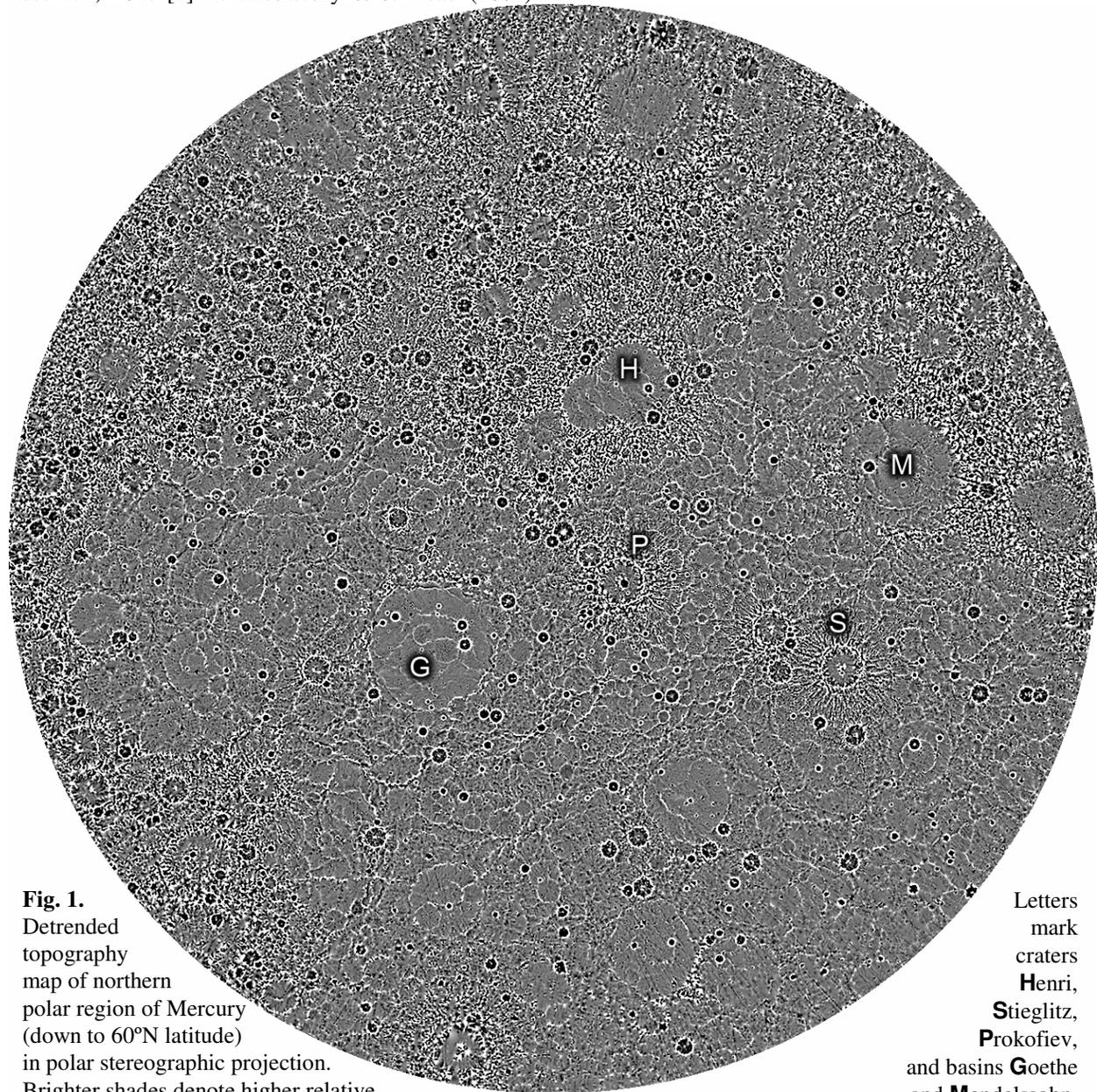


Fig. 1.
Detrended topography map of northern polar region of Mercury (down to 60°N latitude) in polar stereographic projection. Brighter shades denote higher relative elevation.

Letters mark craters **H**enri, **S**tieglitz, **P**rokofiev, and basins **G**oethe and **M**endelssohn.