

### Mn-Cr dating of enstatite chondrites

J. Hopp<sup>1,2</sup>, M. Trierloff<sup>1,2</sup>, J.-C. Storck<sup>1,3</sup>, T. Ludwig<sup>1</sup>, H.-P. Meyer<sup>1</sup>, R. Altherr<sup>1</sup> and A. El Goresy<sup>4</sup>, <sup>1</sup>Institut für Geowissenschaften, Universität Heidelberg, Im Neuenheimer Feld 234-236, D-69120 Heidelberg, jens.hopp@geow.uni-heidelberg.de, <sup>2</sup>Klaus-Tschira-Labor für Kosmochemie, Universität Heidelberg, Im Neuenheimer Feld 234-236, D-69120 Heidelberg, Germany, <sup>3</sup>Institut für Geochemie und Petrologie, ETH Zürich, Sonneggstrasse 5, 8092 Zürich, Switzerland, <sup>4</sup>Bayerisches Geoinstitut, Universität Bayreuth, Universitätsstrasse 30, D-95440 Bayreuth, Germany.

**Introduction:** The  $^{53}\text{Mn}$ - $^{53}\text{Cr}$  system (half-life 3.7 Ma) is a short-lived decay system that can be used as a relative chronometer of early solar system processes. Ideally, if  $^{53}\text{Cr}/^{52}\text{Cr}$ -ratios in a sample are linearly correlated with the associated ratio of (stable)  $^{55}\text{Mn}/^{52}\text{Cr}$  the excess  $^{53}\text{Cr}$  can be interpreted to be the result of in situ decay of  $^{53}\text{Mn}$ . The respective slope of this isochron corresponds to the initial  $^{53}\text{Mn}/^{55}\text{Mn}$  of the sample. This initial  $^{53}\text{Mn}/^{55}\text{Mn}$  ratio can be translated into an absolute age by comparing it with the initial  $^{53}\text{Mn}/^{55}\text{Mn}$  ratio of a standard meteorite of known absolute age. Commonly used is the LEW 86010 angrite Mn-Cr age standard with an absolute U-Pb age of  $4557.8 \pm 0.5$  Ma and an initial  $^{53}\text{Mn}/^{55}\text{Mn}$  ratio of  $(1.25 \pm 0.07) \cdot 10^{-6}$  (errors are  $1\sigma$ -uncertainties) [1].

Mn and Cr are both rather abundant elements and assumed to have a relatively low diffusivity. Application of the Mn-Cr dating method is limited to few carrier phases with a low partitioning behavior of Cr and high Mn/Cr ratios and hence, besides other problematic issues, is still rarely applied in early solar system chronology. Enstatite chondrites contain a large variety of exotic mineral phases formed at low  $f(\text{O}_2)$ , in particular Mn-bearing sulfides, which might be suited for Mn-Cr dating. The aim of our project is to add more cosmochronological constraints on the early evolution of the enstatite chondrite parent bodies.

**Methods:** In this study we screened several distinct mineral phases for potential  $^{53}\text{Cr}$ -excesses in enstatite chondrites at the new Cameca 1280 HR ion probe facility in Heidelberg, Germany. We used a focussed beam mode and the primary beam current commonly was in the range of 2 to 3 nA (equivalent to a spot diameter of about 3-5  $\mu\text{m}$ ). The field aperture was adjusted to values down to about 500 in order to avoid potentially nearby Cr-bearing phases. Element ratios were obtained by electron microprobe analyses with a Cameca SX51 device.

**Samples and results:** We analyzed the enstatite chondrites Sahara 97072 (EH 3), Indarch (EH 4), EET 96135 (EH 4/5), LAP 02225 (EH impact melt), MAC 88136 (EL 3) and Neuschwanstein (EL 6). For some of these chondrites age data from different chronometers already exist (e.g., Mn-Cr [1], [2], [3] I-Xe [4], [5], [6], Ar-Ar [7]). The major Mn-bearing phases in EH and EL chondrites are niningerite and alabandite, respectively. However, Mn/Cr ratios in both phases always were below 400, i.e., far below required ratios to detect  $^{53}\text{Cr}$ -excesses. Hence,  $^{53}\text{Cr}/^{52}\text{Cr}$ -ratios are indistinguishable from the  $^{53}\text{Cr}/^{52}\text{Cr}$ -ratio determined by analyses of Cr-rich phases troilite and daubr elite ( $\text{Mn}/\text{Cr} < 1$ ). Results of troilite and daubr elite analyses were intended for determination of the relative sensitivity factor (rsf) for Mn/Cr (= ratio of Mn/Cr compositions determined by electron microprobe versus ion probe signals). However, the fine-grained nature of daubr elite in most meteorites and the relatively low amount of Mn in troilite made a consistent determination of rsf-values challenging. In case of LAP 02225 a reproducible rsf-value of  $1.43 \pm 0.01$  (two analyses of daubr elite) could be obtained. Only in sphalerite we found significant excess  $^{53}\text{Cr}$  associated with high  $^{55}\text{Mn}/^{52}\text{Cr}$ -ratios (up to  $10^7$ ). However, commonly data do not define a distinct isochron but rather show some scatter. Most sphalerite data were obtained for the impact melt LAP 02225 for which a range in initial  $^{53}\text{Mn}/^{55}\text{Mn}$  of  $(1.5\text{-}3.5) \cdot 10^{-7}$  could be derived, corresponding to an absolute age range of 4546.5-4551.0 Ma (rsf=1.43) which agrees with I-Xe and Ar-Ar ages. This is consistent with fast cooling associated with simultaneous closure of all isotope chronometers, as expected for an impact melt rock. The initial  $^{53}\text{Mn}/^{55}\text{Mn}$  derived for a single sphalerite datum of MAC 88136 is  $(2.55 \pm 0.30) \cdot 10^{-7}$ , equivalent to an absolute age of  $4549.3 \pm 0.6$  Ma (for rsf=1). This is relatively young and may point to a weak disturbance of the Mn-Cr system by a thermal event. For Indarch we obtained an initial  $^{53}\text{Mn}/^{55}\text{Mn} = (4.2 \pm 1.1) \cdot 10^{-6}$  (rsf=1), which corresponds to an absolute age of  $4564.1 \pm 1.5$  Ma, in broad agreement with previous results [1], [2], [3]. Application of a rsf=1.43 as for LAP 02225 would result in 1.9 Ma younger sphalerite ages for both meteorites. A better determination of the Mn/Cr rsf-value for sphalerite is an important aspect of future work. In addition, more Mn-Cr analyses of EL chondrites are planned.

**References:** [1] Shukolyukov A. and Lugmair G. W. 2004. *Geochimica et Cosmochimica Acta* 68:2875-2888. [2] El Goresy et al. 1992. Abstract #1165. 23th Lunar & Planetary Science Conference. [3] Wadhwa M. et al. 1997. *Meteoritics & Planetary Science* 32:281-292. [4] Kennedy B. M. et al. 1988. *Geochimica et Cosmochimica Acta* 52:201-211. [5] Busfield A. et al. 2008. *Meteoritics & Planetary Science* 43:883-897. [6] Hopp J. et al. 2016. *Geochimica et Cosmochimica Acta* 174:196-210. [7] Hopp J. et al. 2014. *Meteoritics & Planetary Science* 49:358-372.