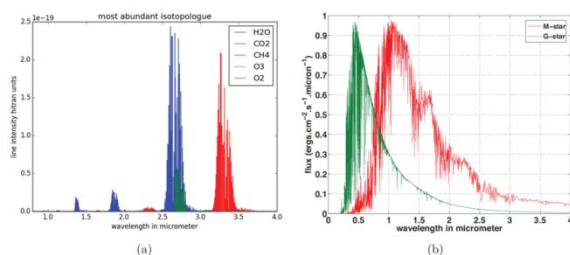


**DETERMINING THE INNER EDGE OF THE HABITABLE ZONE AROUND M-DWARFS USING 3-D CLIMATE MODELS.** R. K. Kopparapu<sup>1,2,3</sup>, E. T. Wolf<sup>4</sup>, J. Haqq-Misra<sup>3,5</sup>, J. F. Kasting<sup>1</sup>, S. Mahadevan<sup>1</sup>, R. Terrien<sup>1</sup>. <sup>1</sup>Pennsylvania State University (ruk15@psu.edu), <sup>2</sup>Virtual Planetary Laboratory, <sup>3</sup>Blue Marble Space Institute of Science, <sup>4</sup>University of Colorado, Boulder ([Eric.Wolf@colorado.edu](mailto:Eric.Wolf@colorado.edu)), <sup>5</sup>jacob@bmsis.org

**Introduction:** The study of habitable zones (HZs) has received significant attention recently with the discoveries of terrestrial mass/size planets from both ground-based surveys and the Kepler mission. Specifically, terrestrial planets near the inner edge of HZs of M-stars have received more scrutiny because of their shorter orbital periods, which increases their chances of detection, and characterization, compared to G-stars. In this presentation, we will report our ongoing efforts to determine the inneredge of the HZ around late-K to late-M-dwarfs, using multiple 3-dimensional global circulation models (GCMs). We are updating the radiative transfer in these GCMs with new absorption coefficients for H<sub>2</sub>O using the most recent line-by-line HITRAN2012 database, and with semi-empirical water-vapor continuum absorption, which will significantly affect the inner edge of the HZ. We will consider both water-rich and dry 'dune'-like planets, to identify the most accurate HZ limits around M-stars.

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**Caption:** *Panel a:* Line intensity of water-vapor (blue) in the near-IR, compared to other molecules. The units of the vertical axis are cm<sup>-1</sup>/(molecule cm<sup>-2</sup>). *Panel b:* Spectrum of an M-dwarf star (red) which emits mostly in the near-IR, where water-vapor is a dominant absorber. H<sub>2</sub>O line intensity from HITRAN2008 database, VPL (<http://vplapps.astro.washington.edu/>). M-star spectrum from BT-Settl grid of models (Allard et al., 2003, 2007).

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